

# LIMIT TREATMENT OF SPACE-TIME FOAM CALCULATIONS IF INITIAL LENGTH OF SPACE-TIME COLLAPSES TO A NON PATHOLOGICAL BIG BANG SINGULARITY ACCESSIBLE TO MODIFIED EINSTEIN EQUATIONS

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## Abstract

When initial radius  $R_{initial} \rightarrow 0$  if Stoica actually derived Einstein equations in a formalism which remove the big bang singularity pathology, then the reason for Planck length no longer holds. We follow what Ng derived as limit calculations as to a space time length factor  $\ell$ . Without  $R_{initial} \rightarrow 0$  the drop off of the vacuum energy as given by  $\Lambda_{Today} \sim \Lambda_{EW} \exp(-H_{EW} \cdot t|_{Today})$  is at least  $10^{-38}$  the value of  $\Lambda_{EW}$ . We review the work by Ng as to quantum foam as to how that effects a general expression as to energy  $\Lambda_{EW}$  when  $R_{initial} \sim \frac{1}{\#} \ell_{Ng} < l_{Planck}$ , with  $\ell_{Ng}$  determined at least approximately by arguments he presented in 2008 in the Dark side of the universe conference. Well before  $\frac{1}{\#} \rightarrow 0^+$  certain effects make themselves apparent, in ways which are illustrated in this manuscript. Having  $\rho \rightarrow \infty$  at a point singularity would remove expansion by the scale factor,  $\rho \sim H^2 / G \Leftrightarrow H \approx a^{-1}$  so that the extreme version of Stoica's treatment in an isolated 4 dimensional universe would be no expansion at all.

Keywords: *Fjortoft theorem, thermodynamic potential, matter creation, vacuum energy Mach's theorem, non pathological singularity affecting Einstein equations, planck length. Braneworlds.*

## 1. Introduction

This article is to investigate what happens physically if there is a non pathological singularity at the start of space-time, i.e. no reason to have a minimum nonzero length. The reasons for such a proposal come from [1] by Stoica who may have removed the reason for the development of Planck's length as a minimum safety net to remove what appears to be unavoidable pathologies at the start of applying the Einstein equations at a space-time singularity. What show are unavoidable collapse of the usual assumptions of the inter relationships of the number of operations in space-time, of the number of bits, and also of the average energy per bit of space time. We will work on the assumption of

$R_{initial} \sim \frac{1}{\#} \ell_{Ng} < l_{Planck}$  and only invoke the  $R_{initial} \rightarrow 0$  value to make an extreme point as to back up assertions made earlier, without calculations in a prior paper. Certain physics effects make themselves apparent well before  $\frac{1}{\#} \rightarrow 0^+$  limit, and are commented upon in this article.

$\rho \sim H^2 / G \Leftrightarrow H \approx a^{-1}$  in particular is remarked upon. This is a counter part to *Fjortoft* theorem in **Appendix I** below.

## 2. Mach's Principle as initially stated, in EW to present day era. Preserving Planck's constant

We first of all review an earlier proposed Mach's principle for the Gravitinos in the electro weak era, and then the 2<sup>nd</sup> modern day Mach's principle, as organized by the author are as seen in [2]. This construction was used in an earlier article to argue in favor of a constant value of  $h$ , i.e. Planck's constant. For the sake of review, we will state that the values in

$$\frac{GM_{electro-weak} \Big|_{Super-partner}}{R_{electro-weak} c^2} \approx \frac{GM_{today} \Big|_{Not-Super-Partner}}{R_0 c^2} \quad (1)$$

are really a statement of information conservation. I.e. the amount of information stored in the left hand side of (1) is the same as information as in the right hand side of (1) above. Here, M as in the electro weak era refers to  $M = N$  times  $m$ , where M is the total 'mass' of the gravitinos, N the number of Gravitinos, and R for the electro weak as an infinitely small spatial radius. Where as the Right hand side is for M for gravitons (not super partner objects) = N as the (number of gravitons) and  $m$  (the ultra low mass of the graviton) in the right hand side of (1) This formular (1) should be compared with a change in entropy formula given by Lee [3] about the inter relationship between energy, entropy and temperature as given by

$$m \cdot c^2 = \Delta E = T_U \cdot \Delta S = \frac{\hbar \cdot a}{2\pi \cdot c \cdot k_B} \cdot \Delta S \quad (2)$$

Lee's formula is crucial for what we will bring up in the latter part of this document. Namely that changes in initial energy could effectively vanish if [1] is right, i.e. Stoica removing the non pathological nature of a big bang singularity.

If the mass  $m$ , i.e. for gravitons is set by acceleration (of the net universe) and a change in entropy  $\Delta S \sim 10^{38}$  between the electroweak regime and the final entropy value of, if  $a \cong \frac{c^2}{\Delta x}$  for acceleration is used, so then we obtain

$$S_{Today} \sim 10^{88} \quad (3)$$

Then we are really forced to look at (1) as a paring between gravitons (today) and gravitinos (electro weak) in the sense of preservation of information.

Having said this, we state that (3) above is based upon certain assumptions usually congruent with the quantum foam model [4], and that what happens, especially to (3) is profoundly affected as we enter a regime for which  $R_{initial} \sim \frac{1}{\#} \ell_{Ng} < l_{Planck}$ . We follow Ng's derivation [4] to make a cautionary point, while removing his worry about black holes, to state something about not only energy  $E$ , but also  $\Lambda_{EW}$  and by extention  $\rho \sim H^2 / G \Leftrightarrow H \approx a^{-1}$ . As  $\rho$  changes due to  $\rho \sim H^2 / G$  and  $R_{initial} \sim \frac{1}{\#} \ell_{Ng} < l_{Planck}$ , then  $a$  is also altered.

The point we make, is to go to the case of having  $\rho \rightarrow \infty$  means there would be no net expansion at all. ***If the universe were 4 dimensional and closed.*** We do not take the case of having no initial energy at the beginning of a closed universe as feasible or even realistic to refer to. The information theory implications though of what Stoica implies [1] as to bits and also holography need to be studied.

What will determine the answer to this question is if  $\Delta E_{initial}$  goes to zero if  $R_{initial} \rightarrow 0$  which happens if there is no minimum distance mandated to avoid the pathology of singularity behavior at the heart of the Einstein equations. In doing this, we avoid using the  $E \rightarrow 0^+$  situation, and instead refer to a nonzero energy, with  $\Delta E_{initial}$  instead vanishing.

### 3. Review of Ng, [4] with comments.

First of all, Ng refers to the Margolus-Levitin theorem with the rate of operations  $< E/\hbar \Rightarrow \#operations < E/\hbar \times time = \frac{Mc^2}{\hbar} \cdot \frac{l}{c}$ . Ng wishes to avoid black-hole formation  $\Rightarrow M \leq \frac{lc^2}{G}$ . This last step is not important to our view point, but we refer to it to keep fidelity to what Ng brought up in his presentation. Later on, Ng refers to the  $\#operations \leq (R_H/l_p)^2 \sim 10^{123}$  with  $R_H$  the Hubble radius. Next Ng refers to the  $\#bits \propto [\#operations]^{3/4}$ . Each bit energy is  $1/R_H$  with  $R_H \sim l_p \cdot 10^{123/2}$

The key point as seen by Ng [4] and the author is in

$$\#bits \sim \left[ \frac{E}{\hbar} \cdot \frac{l}{c} \right]^{3/4} \approx \left[ \frac{Mc^2}{\hbar} \cdot \frac{l}{c} \right]^{3/4} \quad (4)$$

Assuming that E of the universe is not set equal to zero, which the author views as impossible, the above equation says that the number of available bits goes down dramatically if one sets  $R_{initial} \sim \frac{1}{\#} \ell_{Ng} < l_{Planck}$  ? Also Ng writes entropy S as proportional to a particle count via N.

$$S \sim N \cong [R_H / l_p]^2 \quad (5)$$

We rescale  $R_H$  to be

$$R_H|_{rescale} \sim \frac{l_{Ng}}{\#} \cdot 10^{123/2} \quad (6)$$

The upshot is that the entropy, in terms of the number of available particles drops dramatically if # becomes larger.

So, as  $R_{initial} \sim \frac{1}{\#} \ell_{Ng} < l_{Planck}$  grows smaller, as # becomes larger

- a. The initial entropy drops
- b. The number of bits initially available also drops.

The limiting case of (4) and (5) in a closed universe, with no higher dimensional embedding is that both would vanish, i.e. appear to go to zero if # becomes very much larger

#### 4. Examination of Mitra's [5] formation of mass, energy and its possible effects on the cosmological 'contant' vacuum energy.

The prior result was to state that Avesion's [6] time varing  $\hbar(t)$  in fact is a constant value, with no variation as due to alleged behavior represented by Mach's principle as represented by (1) above. What will be done next will be to look at the role of energy of the universe, and what it says about quintessence. The construction comes from Mitra [5] and is adapted to what Beckwith did with the Machian universe relations [1] as given in (1) to (3) above. Mitra [5] in lieu of working with a FRLW universe, wrote

$$E = M(r,t) = \frac{4\pi}{3} \cdot R^3 \rho. \quad (7)$$

$$R = a(t) \cdot r(t)$$

$$E = -\frac{M}{a \cdot r} + \frac{1}{2} \cdot (\dot{a} \cdot r + \dot{r} \cdot a)^2$$

The density factor so parlayed in this treatment in the 1<sup>st</sup> equation in (7) was cited to have the relationship [5] by Mitra

$$\rho \cdot a(t) = const \quad (8)$$

and the author put in, subsequently the following scaling factors

$$\begin{aligned} a(t) &= a_0 \exp(H \cdot t) \\ r(t) &= r_0 \exp(\beta \cdot t) \\ \rho_{\bullet} &= \frac{\Lambda}{8\pi} \end{aligned} \tag{9}$$

In addition is the  $\dot{a} = H \cdot a$  associated with the Hubble parameter and all that This leads to the energy value of the last equation of (7) to be written as

$$(a \cdot r)^3 - \frac{E}{(\beta + H)^2} \cdot (a \cdot r) - \frac{M}{(\beta + H)^2} = 0 \tag{10}$$

Using a typical cubic solution for real valued roots, this comes out to be If we say that E=M, in the sense of the speed of light being set =1, then

$$(a \cdot r) \sim \left[ \frac{M}{(\beta + H)} \right]^{1/3} + H.O.T. \tag{11}$$

This M though is for the total mass of the universe. But still we have

$$a(t) \propto \frac{const}{\rho_{\bullet}} \approx \exp(H \cdot t) \Rightarrow \rho_{\bullet} \propto \Lambda \sim \exp(-H \cdot t) \tag{12}$$

In so many words, the parameter for quintessence goes to almost zero today, i.e.

$$\Lambda \sim \exp(-H \cdot t) \xrightarrow{t \rightarrow \infty} 0^+ \tag{13}$$

Question to ask is as follows. I.e. look at what the author derived

$$(a \cdot r) \sim \left[ \frac{M}{(\beta + H)} \right]^{1/3} \Leftrightarrow M \sim (\beta + H) \cdot (a \cdot r)^3 \tag{14}$$

This equation breaks down if  $R_{initial} \rightarrow 0$ . **What would replace it ?**

- 5. Does it make sense to talk of vacuum energy if  $R_{initial} \neq 0$  is changed to  $R_{initial} \rightarrow 0$ ? Only answerable straightforwardly if an embedding superstructure is assigned. Otherwise difficult.**

The adaptation of the Mitra [5] relation for mass as given by (7) presupposes that there is a well defined nonzero initial radius for cosmological evolution. We summarize what may be the high lights of this inquiry leading to the present paper as follows.

5a. One could have the situation if  $R_{initial} \rightarrow 0$  of an infinite point mass, if there is an initial nonzero energy in the case of just four dimensions and no higher dimensional embedding even if [1] goes through verbatim. The author sees this as unlikely. But is prepared to be wrong. The infinite point mass construction is verbatim if one assumes a closed universe, with no embedding superstructure. Note this appears to nullify the parallel brane world construction author, in lieu of the manuscript sees no reason as to what would perturb this infinite point structure, so as to be able to enter in a big bang era. In such a situation, one would not have vacuum energy.

5b. The most problematic scenario.  $R_{initial} \rightarrow 0$  and no initial cosmological energy. I.e. this in a 4 dimensional closed universe. Then there would be no vacuum energy at all initially. A literal completely empty initial state, which is not held to be viable by Volovik [6].

5c. Finding that additional dimensions are involved, than just 4 dimensions may give credence to the authors speculation as to initial degrees of freedom reaching up to 1000, and the nature of a phase transition from essentially very low degrees of freedom, to over 1000 maybe in fact a chaotic mapping as speculated by the author in 2010 [7].

5d. What the author would be particularly interested in knowing would be if actual semiclassical reasoning could be used to get to an initial prequantum cosmological state. This would be akin to using [8], but even more to the point, using [9] and [10], with both these last references relevant to forming Planck's constant from electromagnetic wave equations. The author points to the enormous Electromagnetic fields in the electroweak era as perhaps being part of the background necessary for such a semiclassical derivation, plus a possible Octonionic space-time regime, as before inflation flattens space-time, as forming a boundary condition for such constructions to occur [11]

The relevant template for examining such questions is given in the following table 1 as printed below.

5e. The meaning of Octonionic geometry prior to the introduction of quantum physics presupposes a form of embedding geometry and in many ways is similar to Penrose's cyclic conformal cosmology speculation Note the following argument, as :

5f. We are stuck with how a semiclassical argument can be used to construct Table 1 below. In particular, we look at how Planck's constant is derived, as in the electroweak regime of space-time, for a total derivative [9],[10]

$$E_y = \frac{\partial A_y}{\partial t} = \omega \cdot A'_y (\omega \cdot (t-x)) \quad (15)$$

Similarly [9],[10]

$$B_z = -\frac{\partial A_y}{\partial x} = \omega \cdot A'_y (\omega \cdot (t-x)) \quad (16)$$

The A field so given would be part of the Maxwell's equations given by [9] as, when [ ] represents a D'Albertain operator, that in a vacuum, one would have for an A field [9], [10]

$$[ ]A = 0 \quad (17)$$

And for a scalar field  $\phi$

$$[ ]\phi = 0 \quad (18)$$

Following this line of thought we then would have an energy density given by, if  $\varepsilon_0$  is the early universe permeability [9]

$$\eta = \frac{\varepsilon_0}{2} \cdot (E_y^2 + B_z^2) = \omega^2 \cdot \varepsilon_0 \cdot A_y'^2(\omega \cdot (t-x)) \quad (19)$$

We integrate (19) over a specified E and M boundary, so that, then we can write the following condition namely [9],[10].

$$\iiint \eta d(t-x) dydz = \omega \varepsilon_0 \iiint A_y'^2(\omega \cdot (t-x)) d(t-x) dydz \quad (20)$$

(20) would be integrated over the boundary regime from the transition from the Octonionic regime of space time, to the non Octonionic regime, assuming an abrupt transition occurs, and we can write, the volume integral as representing [9],[10]

$$E_{gravitational-energy} = \hbar \cdot \omega \quad (21)$$

Our contention for the rest of this paper, is that Mach's principle will be necessary as an information storage container so as to keep the following, i.e. having no variation in the Planck's parameter after its formation from electrodynamics considerations as in (20) and (21). Then by applying [9], [10] we get  $\hbar$  formed by semiclassical reasons and need to have Machs principle (1) to have the same value up to the present era.

$$\hbar(t) \xrightarrow{\text{Apply-Machs-Relations}} \hbar(\text{Constant value}) \quad (22)$$

The question we can ask, is that can we have a prequantum regime commencing for (20) and (21) for  $\hbar$  if  $R_{initial} \sim \frac{1}{\#} \ell_{Ng} < l_{Planck}$ ? And a closed 4 dimensional universe? If so, then what is the necessary geometrial regime of space-time so that the integration performed in (20) can commence properly? Also, what can we say about the formation of (21) above, as a number, # gets larger and larger, effectively leading to  $\frac{1}{\#} \rightarrow 0^+$ ? We need to know seriously how much space-time is needed to form (21) above? Do infinitesimal amounts of space-time suffice in order to fill in the following table as given below? If the answer is no, then when  $\frac{1}{\#} \rightarrow 0^+$

leading to  $R_{initial} \sim \frac{1}{\#} \ell_{Ng} < l_{Planck} \xrightarrow{\# \rightarrow \infty} 0$ , the semiclassical derivation of (20) leading to (21)

may not work. This is the table to consider if  $R_{initial} \sim \frac{1}{\#} \ell_{Ng} < l_{Planck} \xrightarrow{\# \neq \infty} small - value$  and not zero. Also, with an Octonionic geometry regime which is a pre quantum state.[11]

**TABLE 1**

Time Interval	Dynamical consequences	Does QM/WdW apply?
Just before Electroweak era	Form $\hbar$ from early E & M fields, and use Maxwell's Equations with necessary to implement boundary conditions created from change from Octonionic geometry to flat space	NO
Electro-Weak Era	$\hbar$ kept constant due to Machian relations	YES Use (1) as linkage
Post Electro-Weak Era to today	$\hbar$ kept constant due to Machian relations	YES Wave function of Universe

In so many words, the formation period for  $\hbar$  is our pre-quantum regime. This table 1 could even hold if  $R_{initial} \rightarrow 0$  but that the 4 dimensional space-time exhibiting such behavior is embedded in a higher dimensional template

**6. Having not  $R_{initial} \rightarrow 0$ , but  $R_{initial} \sim \frac{1}{\#} \ell_{Ng} < l_{Planck}$  growing smaller, as # becomes larger.**

If there is only an isolated 4 dimensional universe a situation for which.

6a. The initial entropy drops

6b. The number of bits initially available also drops.

Then, we argue that dramatically cutting the initial entropy and also the bits would lead to real trouble as far as preserving the formation of  $\hbar$  as given in [9], [10] and may lead to difficulties in application of (1) especially if bits for a computation for expansion cannot be formed in a general way at the start of inflation.

**7. If  $R_{initial} \rightarrow 0$  then if there is an isolated, closed universe, there is a messy situation.**

One does not have initial entropy, and the number of bits initially disappears.

Abandoning the idea of a completely empty universe, this unperturbed point of matter-energy appears to be a recpede for a static point with no perturbation, as may be the end result of applying Fjortoft theorem [12] to the thermodynamic potential as given in [13], i.e. the non definitive answer for fulfillment of criteria of instability by applying Fjortoft's theorem [12] to the potential [13] leading to no instability as given by the potential given in [13] may lead to a point of space-time with no change, i.e. a singular point with 'infinite' mass which does not change at all.

8. **Can an alternative to a minimum length be put in? Consider the example of Planck time as the minimal component, not Planck length.**

From J. Dickau, the following was given to the author, as a counter part as to how to view thresholds as to how a Mandelbrot set may pre select for critical behavior different from what is being pre supposed in this manuscript.[14]

Dickau writes:

“If we examine the Mandelbrot Set along the Real axis, it informs us about behaviors that also pertain in the Quaternion and Octonic case-because the real axis is invariant over the number types. If numbers larger than .25 are squared and summed recursively ( i.e.  $-z = z^2 + c$  ) the result will blow up, but numbers below this threshold never get to infinity, no matter how many times they are iterated. But once space-like dimensions are added-i.e. an imaginary compoent- the equation blows up exponentially, faser than when iterated“

Dickau concludes:

“Anyhow there may be a minimum (space-time length) involved but it is probably in the time direction”.

This is a counter pose to the idea of minimum length, i.e. the idea being a replacement for what the author put in here: looking at a beginning situation with a crucial parameter  $R_{initial}$  even if the initial time step is “put in by hand”. First of all, look at [4], if E is M, due to setting  $c = 1$ , then

$$\Delta E_{initial} \approx 4\pi\rho(R_{initial})^2 \Delta R_{initial} \quad (23)$$

Everything depends upon the parameter  $R_{initial}$  which can go to zero. The choice as to  $R_{initial}$  going to zero, or not going to zero will be conclusion of our article.

We have to look at what (23) tells us, even if we have an initial time step for which time is initially indeterminate, as given by a redoing of Mitra’s  $g_{00}$  formula [4] which we put in to establish the indeterminacy of the initial time step if quantum processes hold.

$$\left( g_{00} = \exp \left[ \frac{-2}{1 + (\rho(t) / p(t))} \right] \right) \xrightarrow{\rho+p=0} 0 \quad (24)$$

What Dickau is promoting is, that the Mandelbrot set, if applicable to early universe geometry, that what the author wrote, with  $R_{initial} \sim \frac{1}{\#} \ell_{Ng} < l_{Planck} \xrightarrow{\# \neq \infty} small - value$  potentially going to zero, is less important than a minimum time length. To which the author states, if Dickau is correct as to applicability of the Mandelbrot set, that he, the author is happily corrected, but he also thinks that the Mandelbrot set is a beautiful example of the fungability of space – time metrics used. I.e. how one sets the initial space-time potential is to determine the correctness of the Mandelbrot set. I.e. the [12] reference, as given, by Thanu Padmanabhan appears not to have a Mandelbrot set, in its thermodynamic potential. The instability issue is

reviewed in **Appendix II**. for those who are interested in the author's views as to lack proof of instability. It uses [12] which the author views as THE reference as far as thermodynamic potentials and the early universe.

## 7. We need to re consider the role of Quantum gravity models at the onset of inflation.

*We are stuck in all Quantum gravity models as of putting in an initial time step 'by hand' so to speak which raises fundamental issues of what would form an initial time step in Quantum gravity. How the transition from the left to the right hand side of (22) occurs is crucial and it comes about because of a transition from Octonionic geometry to quantum accessible and analyzable flat space geometry.*

**7a. Having not**  $R_{initial} \rightarrow 0$ , but  $R_{initial} \sim \frac{1}{\#} \ell_{Ng} < l_{Planck}$  grows smaller, as # becomes

larger leads to , **if there is only an isolated 4 dimensional universe** a situation for which.

- i. The initial entropy drops
- ii The number of bits initially available also drops.

We argue that dramatically cutting the initial entropy and also the bits would lead to real trouble as far as preserving the invariance of  $\hbar$  [15]

**7b.** If  $R_{initial} \rightarrow 0$  then if there is an isolated, closed universe, one does not have initial entropy, and the number of bits initially disappears. i.e. in lieu of [12] and [13] there may be no perturbation from an infinite point of space-time which remains invariant.

## 8. Conclusions

**8a. The universe if**  $R_{initial} \rightarrow 0$  [1] and if it is an isolated system, i.e. not as embedded in higher dimensions as referred to in [16] may have no bits, or computations as thought of by Ng [4]. This would be in tandem with the authors conclusion that one would have an initial infinite point mass and no evolution. And no generation of entropy.

**8b. If**  $R_{initial} \rightarrow 0$  [1] but the universe is embedded in a higher dimensional system, as given by [16], then there is no reason to say there are no bits, or computations, and the universe will continue to evolve with entropy as a by product of that evolution.

**8c. The universe if one does not have**  $R_{initial} \rightarrow 0$ , i.e. if [1] does not hold, then there will be bits, entropy being generated, and also loop quantum gravity and quantum measures [17] so long as the minimum  $l_p$  Planck length grid exists. The reacceleration of the universe commences as given in [18] due to DE being the same as vacuum energy. The author in **Appendix III** gives a derivation as to how the DE will be put in, via branes, which holds for both 8b and 8c , as a to be proven sideline.

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### Appendix I. Fjortoft theorem:

A **necessary condition for instability** is that if  $z_*$  is a point in spacetime for which  $\frac{d^2U}{dz^2} = 0$  for any given potential  $U$ , then there must be some value  $z_0$  in the range  $z_1 < z_0 < z_2$  such that

$$\left. \frac{d^2U}{dz^2} \right|_{z_0} \cdot [U(z_0) - U(z_*)] < 0 \quad (1)$$

For the proof, see [11] and also consider that the main discussion is to find instability in a physical system which will be described by a given potential  $U$ . Next, we will construct in the boundary of the EW era, a way to come up with an optimal description for  $U$

### Appendix II. Constructing an appropriate potential for using Fjortoft theorem in cosmology for the early universe cannot be done. We show why

To do this, we will look at Padamanabhan [12] and his construction of (in Dice 2010) of thermodynamic potentials he used to have another construction of the Einstein GR equations. To start, Padamanabhan [5] wrote

If  $P_{cd}^{ab}$  is a so called Lovelock entropy tensor, and  $T_{ab}$  a stress energy tensor

$$\begin{aligned}
U(\eta^a) &= -4 \cdot P_{ab}^{cd} \nabla_c \eta^a \nabla_d \eta^b + T_{ab} \eta^a \eta^b + \lambda(x) g_{ab} \eta^a \eta^b \\
&= U_{gravity}(\eta^a) + U_{matter}(\eta^a) + \lambda(x) g_{ab} \eta^a \eta^b \\
&\Leftrightarrow U_{matter}(\eta^a) = T_{ab} \eta^a \eta^b; U_{gravity}(\eta^a) = -4 \cdot P_{ab}^{cd} \nabla_c \eta^a \nabla_d \eta^b
\end{aligned} \tag{1}$$

We now will look at

$$\begin{aligned}
U_{matter}(\eta^a) &= T_{ab} \eta^a \eta^b \quad ; \\
U_{gravity}(\eta^a) &= -4 \cdot P_{ab}^{cd} \nabla_c \eta^a \nabla_d \eta^b
\end{aligned} \tag{2}$$

So happens that in terms of looking at the partial derivative of the top (1) equation, we are looking at

$$\frac{\partial^2 U}{\partial (\eta^a)^2} = T_{aa} + \lambda(x) g_{aa} \tag{3}$$

Thus, we then will be looking at if there is a specified  $\eta_*^a$  for which the following holds.

$$\left[ \frac{\partial^2 U}{\partial (\eta^a)^2} = T_{aa} + \lambda(x) g_{aa} \right]_{\eta_0^a}^* < 0 \tag{4}$$

$$\left[ -4 \cdot P_{ab}^{cd} (\nabla_c \eta_0^a \nabla_d \eta_0^b - \nabla_c \eta_*^a \nabla_d \eta_*^b) + T_{ab} \cdot [\eta_0^a \eta_0^b - \eta_*^a \eta_*^b] + \lambda(x) g_{ab} \cdot [\eta_0^a \eta_0^b - \eta_*^a \eta_*^b] \right]$$

What this is saying is that there is no unique point, using this  $\eta_*^a$  for which (4) holds. Therefore, we say there is no official point of **instability of  $\eta_*^a$**  due to (3). The Lagrangian structure of what can be built up by the potentials given in (3) with respect to  $\eta_*^a$  mean that we cannot expect an inflection point with respect to a 2<sup>nd</sup> derivative of a potential system. Such an inflection point designating a speed up of acceleration due to DE exists a billion years ago [18]. Also note that the reason for the failure for (4) to be congruent to Fjoroft's theorem is due to

$$\left[ \frac{\partial^2 U}{\partial (\eta^a)^2} = T_{aa} + \lambda(x) g_{aa} \right] \neq 0, \text{ for } \forall \eta_*^a \text{ choices} \tag{5}$$

### Appendix III. Modifying a parallel brane-anti brane argument to obtain Massive gravitons

**Part I, of Appendix III. We come into this situation if , as in applying Fjorofts theorem, that we find Appendix II, equation (5), holds, and then if so, we have to find another way to induce vacuum energy.**

What (5), appendix II tells us is that there is an embedding structure for early universe geometry, some of which may take the form of the following diagram.

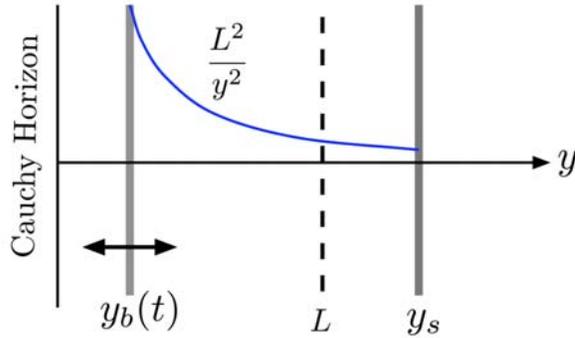


Figure 1, from [16]

**Part II of Appendix III.. Working with a way to achieve energy injection into the universe, without appealing to Fjortoft theorem for alleged instabilities starting from Padmanabhan thermodynamic potential terms**

Padmanabhan [12] introduced the following discussion as to entropy, namely starting with energy, we have

$$E = \frac{1}{2} k_B \int dn T_{loc} \quad (1)$$

And the n value as in (7) is given by

$$dn = 32\pi \cdot P_{cd}^{ab} \cdot \varepsilon_{ab} \cdot \varepsilon^{cd} \cdot dA \quad (2)$$

Where  $P_{cd}^{ab}$  is a so called Lovelock entropy tensor, and  $\varepsilon_{ab}$  a bi normal on the co dimension -2 cross section, and then entropy is stated to be

$$S \propto \int_{\partial v} dn \propto \int_{\partial v} 32\pi \cdot P_{cd}^{ab} \cdot \varepsilon_{ab} \cdot \varepsilon^{cd} \sqrt{\sigma} d^{D-2} x \quad (3)$$

The end result, is that energy is induced via the temperature  $T_{loc}$ , while [12]

$$T_{loc} = N \frac{a^\mu n_\mu}{2\pi} = \text{local acceleration temperature} \quad (4)$$

Also, the change in n can be given by, if  $l_p$  is the Planck's length value[5]

$$\Delta n = \sqrt{\sigma} d^2 x / l_p \quad (5)$$

Looking at (9) and (11) we state that the change in number count given in (4) is really a holographic surface pheonmena, with N defined [12]

$$N = E / [(1/2) k_B T] \quad (6)$$

The upshot is that we can, as implied by Ng[4] easily reference a change in entropy via[4]

$$S \sim n \quad (7)$$

While having a change in n as due to a change in the spatial surface of spacetime as given in (6), we have to realistically infer that the local acceleration temperature (4) is from another pre universe construction and that **local instability is ruled out by Appendix II, equation (5)**. This leads us to ask as to what would be an acceptable way to form the formation of mass, i.e. say the

mass of a graviton, via external factors introduced into our universe prior to the Electroweak era, in cosmology. To do that, look at if there are two branes on the  $AdS_5$  space-time so that with one moving and one stationary, we can look at Figure 1 as background as to introduce such external factors in our present space-time universe during its initial expansion phase

**Part III, of Appendix III: Fall out from adopting Figure 1 and that due to no instability in the Padamanabhan supplied potentials. i.e. a way to obtain graviton mass via a root finding method.**

Using [16] what we find is that there are two branes on the  $AdS_5$  space-time so that with one moving and one stationary, we can look at figure 1 which is part of the geometry used in the spatial decomposition of the differential operator acting upon the  $h_{\bullet}$ . Fourier modes of the  $h_{ij}$  operator [16]. As given by [16], we have that

$$\left[ \partial_t^2 + k^2 - \partial_y^2 + \frac{3}{y} \cdot \partial_y \right] h_{\bullet} = 0 \quad (8)$$

Using [16] the solution to (14) above takes the form of having

$$h_{\bullet} = H_{ij} = e_{ij} \cdot \exp[i \cdot \omega \cdot t] \cdot (m \cdot y)^2 \cdot A \cdot J_2(m \cdot y) \quad (9)$$

$e_{ij}$  is a polarization tensor, and the function  $J_2(my)$  is a 2<sup>nd</sup> order Bessel function[19]. A generalization offered by Durrer et al. [16] leads to

$$h = \left\{ \exp[i \cdot \omega \cdot t] \cdot (m \cdot y)^2 \cdot A \cdot J_2(m \cdot y) \right\} \cdot \left( 1 + \frac{\pi}{4} \cdot (m \cdot \ell)^2 \right) \quad (10)$$

With the factor of  $\left( 1 + \frac{\pi}{4} \cdot (m \cdot \ell)^2 \right)$  coming in due to a boundary condition upon the wall of a brane put in, i.e. looking at [16]. With the right hand side of (10) due to a domain wall tension of a brane.

$$-2 \cdot \partial_y H_{ij} = \kappa_5 \cdot \pi_{ij}^{(T)} \rightarrow 0 \quad (11)$$

This will be in our example set as not equal to zero, in the right hand side, but equal to an extremely small parameter, namely

$$\partial_y H_{ij} \Big|_{y=yb} = \kappa_5 \cdot \pi_{ij}^{(T)} \sim \xi^+ \quad (12)$$

With this turned into

$$\partial_y h \Big|_{y=yb} \sim \delta^+ \quad (13)$$

The right hand side of (13) represents very small brane tension, which is understandable. Then using [16],[19], i.e.

$$\partial_y h \Big|_{y=yb} = \partial_y \left\{ \exp[i\omega t] \cdot (my)^2 \cdot A \cdot J_2(my) \right\} \cdot \left( 1 + \frac{\pi}{4} \cdot (m \cdot \ell)^2 \right) \Big|_{y=yb} \sim \delta^+ \quad (14)$$

and

$$J_2(my) = \frac{(my)^2}{2^2 \cdot 2!} \cdot \left( 1 - \frac{(my)^2}{2^2 \cdot 3} + \frac{(my)^4}{2^4 \cdot 2! \cdot 3 \cdot 4} - \frac{(my)^6}{2^6 \cdot 4! \cdot 3 \cdot 4 \cdot 5} + \dots \right) \quad (15)$$

The upshot is, that afterwards,

$$\begin{aligned} & \frac{(my)^4}{2^2 \cdot 2!} \cdot \frac{1}{y} \left[ \begin{aligned} & \left( 1 - \frac{(my)^2}{2^2 \cdot 3} + \frac{(my)^4}{2^4 \cdot 2! \cdot 3 \cdot 4} - \frac{(my)^6}{2^6 \cdot 4! \cdot 3 \cdot 4 \cdot 5} + \dots \right) \\ & - \left( \frac{2 \cdot (my)^2}{2^2 \cdot 3} + \frac{4 \cdot (my)^4}{2^4 \cdot 2! \cdot 3 \cdot 4} - \frac{6 \cdot (my)^6}{2^6 \cdot 4! \cdot 3 \cdot 4 \cdot 5} + \dots \right) \end{aligned} \right] \\ & = \frac{\delta^+ \cdot \exp[\mp i\omega t]}{A} \cdot \left[ 1 - \frac{\pi}{4} \cdot (m \cdot \ell)^2 \right] \end{aligned} \quad (16)$$

Should the term

$$\frac{\delta^+ \cdot \exp[\mp i\omega t]}{A} \cdot \left[ 1 - \frac{\pi}{4} \cdot (m \cdot \ell)^2 \right] \xrightarrow{\delta^+ \rightarrow 0} 0 \quad (17)$$

Then, (17) is acting much as in [16], whereas, one is recovering a simple numerical exercise as to obtain a suitable solution as given by (18), and (19) due to [3] where the domain tension of the brane vanishes. The novelty as to this approach given in (17) is to obtain a time dependent behavior of the mass of the graviton,

$$(my) = f(t) \Leftrightarrow m \equiv \frac{f(t)}{y} \quad (18)$$

Needless to say, (16) can only be solved for, numerically, i.e. fourth order polynomial solutions for quartic equations still give over simplified dynamics, especially if (18) holds, and makes things more complicated. This is all being done to keep fidelity with respect to [16] as a possible feature of brane world dynamics as reflected in [16], as well as certain issues brought up in [8], [20] as to what is a semiclassical argument can obtain a usually quantum result. If this semiclassical result is true, it has profound implications for [21], and [22] which are held to be M theory results with no classical analogues.